**Herschel-ATLAS: Blazars in the science demonstration phase field**


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**Abstract**

To investigate the poorly constrained sub-mm counts and spectral properties of blazars we searched for these in the **Herschel-ATLAS** (H-ATLAS) science demonstration phase (SDP) survey catalog. We cross-matched 500 µm sources brighter than 50 mJy with the FIRST radio catalogue. We found two blazars, both previously known. Our study is among the first blind blazar searches at sub-mm wavelengths, i.e., in the spectral regime where little is still known about the blazar SEDs, but where the synchrotron peak of the most luminous blazars is expected to occur. Our early results are consistent with educated extrapolations of lower frequency counts and question indications of substantial spectral curvature downwar ds and of spectral upturns at mm wavelengths. One of the two blazars is identified with a Fermi/LAT γ-ray source and a WMAP source. The physical parameters of the two blazars are briefly discussed. These observations demonstrate that the H-ATLAS survey will provide key information about the physics of blazars and their contribution to sub-mm counts.

**Key words.** BL Lacertae objects: general – quasars: general – submillimeter: general

1. Introduction

Blazars, comprising BL Lac objects and flat-spectrum radio quasars, are a class of active galactic nuclei characterized by remarkable properties: high luminosities at the two extreme ends of the electromagnetic spectrum (i.e., in the radio and γ-ray bands), rapid variability, apparent superluminal jet speeds, flat or inverted radio spectrum, and high polarization in the optical waveband. They are interpreted in terms of relativistic effects due to jets propagating at velocities close to the speed of light almost along the line-of-sight. Their spectral energy distribution (SED) can be accounted for as a combination of synchrotron emission, peaking at a frequency $ν_p$ that varies from $10^{12}$ Hz to $10^{19}$ Hz (Nieppola et al. 2006), and inverse Compton scattering, assumed to be responsible for the γ-ray peak.

Padovani & Giommi (1995) classified BL Lac objects into 3 subclasses (low, intermediate, and high frequency synchrotron peaked BL Lacs) depending on the value of $ν_p$. Abdo et al. (2010a) extended the classification to all blazars, subdivided into low (LSP; $ν_p < 10^{14}$ Hz), intermediate (ISP; $10^{14} < ν_p < 10^{15}$ Hz), and high (HSP; $ν_p > 10^{15}$ Hz) synchrotron peaked blazars. For many years, the blazar selection has relied mostly on either radio or X-ray surveys. Radio surveys preferentially select LSPs, while X-ray surveys favor the HSPs. The resulting global distribution of $ν_p$ is thus bimodal, although objects peaking at intermediate frequencies have also been found. The high sensitivity and nearly uniform sky coverage of the Fermi Gamma-ray Space Telescope made it a powerful tool for providing large blazar samples (Abdo et al. 2010b).

In the present paper, we report on the blazars selected at sub-mm wavelengths in the 16 deg$^2$ **Herschel-ATLAS** (H-ATLAS, Eales et al. 2010) science demonstration phase (SDP) field. The **Herschel** data cover the poorly explored frequency range close to the synchrotron peak of the most luminous LSPs (Fossati et al. 1998; Padovani et al. 2006); $ν_p$ carries crucial information about key physical parameters, such as the Lorentz factor $γ_p$ of emitting electrons, the Doppler factor $δ$, and the magnetic field strength $B (v_p ∝ γ_p δ B)$.

Evolutionary models (De Zotti et al. 2005) predict $0.15$ blazars deg$^{-2}$ brighter than 50 mJy at 500 µm, the approximate 5σ detection limit of the H-ATLAS survey (Rigby et al., in prep.), so that the full survey may be expected to yield a sample of $\approx 80$ blazars selected, for the first time, at sub-mm wavelengths. This prediction assumes a flat radio spectral index $α (α ≈ 0.1; S_ν ∝ ν^{-α})$ followed by a parabolic decline close to and above a synchrotron peak frequency that increases with decreasing radio luminosity (Fossati et al. 2008). However, spectral measurements have provided evidence of a substantial spectral curvature downwards from cm to mm wavelengths.
Table 1. The two H-ATLAS blazars found by cross-matching with the FIRST catalog.

<table>
<thead>
<tr>
<th>H-ATLAS name</th>
<th>Δ (arcsec)</th>
<th>S_{1.4 GHz} (mJy)</th>
<th>S_{500 μm} (mJy)</th>
<th>S_{550 μm} (mJy)</th>
<th>S_{250 μm} (mJy)</th>
<th>S_{170 μm} (mJy)</th>
<th>S_{100 μm} (mJy)</th>
</tr>
</thead>
<tbody>
<tr>
<td>J090910.1+012135</td>
<td>0.3</td>
<td>571.7</td>
<td>265.8 ± 8.9</td>
<td>193.8 ± 7.2</td>
<td>159.6 ± 6.4</td>
<td>105.4 ± 26.5</td>
<td>137.3 ± 26.5</td>
</tr>
<tr>
<td>J090940.3+015957</td>
<td>10.3</td>
<td>317.5</td>
<td>68.2 ± 9.1</td>
<td>55.0 ± 7.2</td>
<td>35.5 ± 6.4</td>
<td>&lt;76.2</td>
<td>&lt;79.5</td>
</tr>
</tbody>
</table>

Notes. Coordinates (RA and Dec) refer to the H-ATLAS source. Δ (arcsec) is the angular separation between FIRST and Herschel positions. The quoted errors in Herschel fluxes are statistical only. Calibration errors are estimated to be ~15% for SPIRE (Swinyard et al. 2010), ~10% at 100 μm, and ~20% at 160 μm. Upper limits are at the 3σ level.

In this paper, we describe (Sect. 2) the identification of blazars among sources detected in the H-ATLAS SDP field and discuss (Sect. 3) the properties of the 2 blazars found. In Sect. 4 we summarize our main conclusions.

2. Blazar candidates

Observations were performed with the ESA Herschel Space Observatory (Pilbratt et al. 2010). The 4deg × 4deg H-ATLAS SDP field, centered on (09h05m30s, +00°30′00″), was observed with the Spectral and Photometric Imaging REceiver (SPIRE, Griffin et al. 2010) at 250, 350, and 500 μm and with the Photodetector Array Camera and Spectrometer (PACS, Poglitsch et al. 2010) at 100 and 160 μm. Details of the SPIRE and PACS mapmaking are given by Pascale et al. (in prep.) and Ibar et al. (in prep.), respectively. Sources were extracted from the 250 μm map (see Rigby et al., in prep.). For each source, 350 and 500 μm flux densities were estimated by extracting flux from the appropriate noise-weighted beam-convolved map at the source position determined in the 250 μm map. Fluxes at 100 and 160 μm were then assigned by matching to ≥3σ PACS sources within a positional tolerance of 10″.

The selection of blazar candidates among H-ATLAS sources brighter than 50 mJy was made by identifying FIRST counterparts within a search radius of 15″, the σ_p = FWHM/2.35 of a Gaussian model of the 500 μm beam. The probability that a true counterpart has an apparent positional offset ≥Δ is

\[ p(Δ) = \exp[-0.5(\Delta/σ_p)^2], \]

where σ_p is the global positional uncertainty, taken into account both H-ATLAS and FIRST positional errors. Sources detected with signal-to-noise ratio (S/N) greater than 10 at 250 μm, such as the blazer [HB89] 0906+015 (J090910.1+012135; see Table 1), have a σ_{H-ATLAS} = 1.4″. For PKS 0907+022 (J090940.3+015957), which has data of lower S/N at 250 μm, σ_{H-ATLAS} = 4.9″. The FIRST positions are accurate to better than 1 arcsec at the survey limit and to better than 0.5 arcsec for sources brighter than 5 mJy at 1.4 GHz (White et al. 1997).

With our choice of the search radius, the probabilities of having missed true associations at separations Δ > 15″ are ~8 × 10^{-10} and ~10^{-5}, respectively, for the two values of σ_{H-ATLAS} and σ_{FIRST} = 0.5″.

We note that the two values of σ_{H-ATLAS} are representative of sources with data of high and low S/N at 250 μm.

Since the distribution of FIRST sources is to a large degree uniform, the probability that one of them lies by chance within the angular radius Δ from a given Herschel source is

\[ n_c = \pi Δ^2 n_{1.4 GHz}(S), \]

where S is the flux of the sources we match to. The FIRST catalog is complete down to 0.75 mJy/beam. The corresponding surface density is n_{1.4 GHz}(>0.75 mJy) ≃ 89 deg^{-2}, implying that the maximum probability of a chance association within 15 arcsec is 5 × 10^{-5}. Since there are 185 sources with S_{500 μm} > 50 mJy in the H-ATLAS SD field, the expected number of chance associations is ~0.9. Simulations made by offsetting the radio positions at random and re-running the cross-matching yielded n_c = 0.9(+2.4,-0.8) (68% confidence errors), consistent with expectations for Poisson statistics.

The cross-matching yielded 19 matches. Of them, 10 have far-IR/sub-mm colours typical of low-z dusty galaxies. The far-IR/sub-mm SEDs of 4 others are consistent with those of dusty galaxies at z = 1−2 and their radio emission obeys the radio-FIR correlation (Yun et al. 2001). The other three sources are more than 4σ (σ being the combined rms positional uncertainty) away from their nearest FIRST source, and are therefore very unlikely associations. We note that having 3 spurious associations within our search radius is fully consistent with the results of our simulations. All of these 17 sources were discarded. We note that we aim to select only sources with non-thermal emission brighter than 50 mJy at 500 μm, so we are not interested here in the weaker blazar nuclei that may be hiding in dusty galaxies. The FIRST and Herschel photometric data for the 2 remaining sources are shown in Table 1.

Both H-ATLAS J090910.1+012135 ([HB89] 0906+015) and H-ATLAS J090940.3+015957 (PKS 0907+022) are known blazars. The former is a flat-spectrum quasar with strong broad emission lines at a measured redshift of 1.018 (Falomo et al. 1994). The latter is a BL Lac with an estimated photometric redshift of 1.575 (Richard et al. 2004); a Keck spectrum obtained in 1994. The latter is a BL Lac with an estimated photometric redshift of 1.575 (Richard et al. 2004); a Keck spectrum obtained in 1994. The latter is a BL Lac with an estimated photometric redshift of 1.575 (Richard et al. 2004); a Keck spectrum obtained in 1994. The latter is a BL Lac with an estimated photometric redshift of 1.575 (Richard et al. 2004); a Keck spectrum obtained in 1994. The latter is a BL Lac with an estimated photometric redshift of 1.575 (Richard et al. 2004); a Keck spectrum obtained in 1994.

3. Spectral energy distributions of [HB89] 0906+015 and PKS 0907+022

The photometric data for the two previously known blazars are shown in Figs. 1 and 2. To the data available from NED, we

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added the H-ATLAS fluxes, the UKIDSS data (Lawrence et al., 2007), new ATCA measurements around 20 GHz, new GMRT measurements at 332 MHz (Jarvis et al., in prep.), new Spitzer data (Jarvis et al., in prep.) for [HB89] 0906+015 along with 3σ upper limits that we obtained from IRAS maps (see Table 2). The ATCA measurements were performed by Massardi and Bonavera, in the framework of the Planck-ATCA coeval observation (PACO) project.

The blazar [HB89] 0906+015 was detected by the WMAP satellite (source 215 in the NEWPS5yr catalog; Massardi et al. 2009) and also by the Fermi/Large Area Telescope (LAT) (Abdo et al. 2010b; note that in the previous bright gamma-Ray source list the γ-ray source was identified with PKS 0907+022). The shape of its SED exhibits the typical double hump distribution, attributable to synchrotron and inverse Compton, respectively.

To determine the SEDs, we followed the approach of Abdo et al. (2010a). First we fitted the part of the SEDs dominated by the synchrotron emission using a third degree polynomial. From these fits, we obtained an estimate of the synchrotron peak energy, $\nu^p$, and peak intensity, $S^p(\nu^p)$. Since [HB89] 0906+015 has strong broad emission lines, its optical-UV flux is most likely dominated by the thermal emission produced by the accretion disk, with only a minor contribution from the beamed non-thermal continuum (Ghisellini et al. 2010b). There are indeed clear indications of an optical-UV bump. To avoid biasing estimates of the synchrotron peak frequency and intensity, we excluded these data from the fit. Abdo et al. (2010a) demonstrated that the peak of the inverse Compton SED component, $\nu^p_{IC}$, is strongly correlated with the γ-ray photon index ($\Gamma$). We used their best fit $\nu^p_{IC} - \Gamma$ relationship (their Eq. (5)) to estimate $\nu^p_{IC}$ for the HB89 blazar ($\Gamma = 2.74$). Finally we obtained the peak intensity by fitting the X-ray to γ-ray data points with a parabola peaking at $\nu^p_{IC}$ in the log($S(\nu)/\nu$)–log $\nu$ plane.

The SED parameters are summarized in Table 3. For PKS 0907+022, since no X-ray nor γ-ray measurements are available, we could only use the data fitting method for the synchrotron peak. In both cases, the available data from different instruments are non-simultaneous. As for all blazars, they may therefore be affected by variability, which causes additional uncertainty in the derived spectral parameters. This adds to the uncertainties in estimates of spectral parameters. Based on the best-fit model values of $\nu^p$, both blazars can be classified as LSPs; however, an ISP classification cannot be ruled out.

4. Conclusions

The Herschel detection of two blazars in the SDP area is interesting in a number of ways. First of all, it confirms that the H-ATLAS survey can play an important role in determining properties of blazars in a crucially important spectral region. With an area of 550 deg², the full H-ATLAS should yield a sample of ≈80 blazars, the first statistically significant sample selected at sub-mm wavelengths. The large area is of crucial importance in this context, because the blazar counts are expected to be relatively flat. The H-ATLAS survey will extend by about one order of magnitude downwards in flux the blazar number counts from the PLANCK survey, which is predicted to have a (95% reliability) detection limit of 430 mJy at 545 GHz (Leach et al. 2008).

### Table 2. GMRT, ATCA, UKIDSS, and Spitzer flux density measurements and IRAS 3σ upper limits for the two blazars.

<table>
<thead>
<tr>
<th>Name</th>
<th>[HB89] 0906+015</th>
<th>PKS 0907+022</th>
</tr>
</thead>
<tbody>
<tr>
<td>$S_{332,\text{MHz}}$</td>
<td>941</td>
<td>144</td>
</tr>
<tr>
<td>$S_{18,\text{GHz}}$</td>
<td>1101±55</td>
<td>97±5</td>
</tr>
<tr>
<td>$S_{24,\text{GHz}}$</td>
<td>742±37</td>
<td>69±4</td>
</tr>
<tr>
<td>$S_{1.03,\mu\text{m}}$</td>
<td>0.38</td>
<td>0.22</td>
</tr>
<tr>
<td>$S_{1.25,\mu\text{m}}$</td>
<td>0.97</td>
<td>0.24</td>
</tr>
<tr>
<td>$S_{1.65,\mu\text{m}}$</td>
<td>0.47</td>
<td>0.34</td>
</tr>
<tr>
<td>$S_{2.20,\mu\text{m}}$</td>
<td>0.80</td>
<td>0.54</td>
</tr>
<tr>
<td>$S_{3.6,\mu\text{m}}$</td>
<td>–</td>
<td>–</td>
</tr>
<tr>
<td>$S_{4.5,\mu\text{m}}$</td>
<td>2.05±0.10</td>
<td>–</td>
</tr>
<tr>
<td>$S_{5.8,\mu\text{m}}$</td>
<td>2.80±0.12</td>
<td>–</td>
</tr>
<tr>
<td>$S_{8.0,\mu\text{m}}$</td>
<td>4.2±0.2</td>
<td>–</td>
</tr>
<tr>
<td>$S_{24,\mu\text{m}}$</td>
<td>12.1±0.3</td>
<td>–</td>
</tr>
<tr>
<td>$S_{12,\mu\text{m}}$</td>
<td>&lt;23</td>
<td>&lt;38</td>
</tr>
<tr>
<td>$S_{25,\mu\text{m}}$</td>
<td>&lt;42</td>
<td>&lt;58</td>
</tr>
<tr>
<td>$S_{60,\mu\text{m}}$</td>
<td>&lt;40</td>
<td>&lt;38</td>
</tr>
<tr>
<td>$S_{100,\mu\text{m}}$</td>
<td>&lt;84</td>
<td>&lt;145</td>
</tr>
</tbody>
</table>

### Table 3. SED parameters for the two blazars. Peak fluxes, $\nu F(\nu)$, in erg cm⁻² s⁻¹.

<table>
<thead>
<tr>
<th>Name</th>
<th>[HB89] 0906+015</th>
<th>PKS 0907+022</th>
</tr>
</thead>
<tbody>
<tr>
<td>log($\nu^p$)[Hz]</td>
<td>12.6</td>
<td>13.3</td>
</tr>
<tr>
<td>log($S^p(\nu^p)$)</td>
<td>–11.5</td>
<td>–12.0</td>
</tr>
<tr>
<td>log($\nu^p_{IC}$)[Hz]</td>
<td>20.9</td>
<td>–</td>
</tr>
<tr>
<td>log($S^p_{IC}(\nu^p_{IC})$)</td>
<td>–10.3</td>
<td>–</td>
</tr>
</tbody>
</table>
The number of detected blazars is consistent with the predictions of the De Zotti et al. (2005) model which yields 0.15 blazars deg$^{-2}$ brighter than 50 mJy at 500 $\mu$m, i.e., 2.4 objects in the 16 deg$^{-2}$ SDP area. An observational estimate of the blazar surface density at 5 GHz was obtained by Padovani et al. (2007, see their Figs. 4 and 6). Taking into account the steepening at frequencies approaching the synchrotron peak, the effective 5 GHz to 500 $\mu$m spectral index yielded by the model is $\alpha = 0.2$. The 500 $\mu$m flux limit of 50 mJy then corresponds to 130 mJy at 5 GHz; our results are then in very good agreement with the corresponding blazar surface density found by Padovani et al. (2007) and the relatively flat spectral index assumed by the de Zotti et al. model. If this result is confirmed by data of the full H-ATLAS blazar sample, future models will need to explain in addition evidence of a substantial spectral curvature downwards from cm to mm wavelengths (Waldrum et al. 2007; Sadler et al. 2008).

The blazar catalog of Massaro et al. (2009) lists three more blazars in the SDP field: PKS0858-004 with $S_{1.4\,\text{GHz}} = 325$ mJy, 1RXS J085920.6+0047 with $S_{1.4\,\text{GHz}} = 39$ mJy, and 1RXS J085749.8+0135 with $S_{1.4\,\text{GHz}} = 89$ mJy. None of them has a plausible counterpart among the H-ATLAS sources. In the case of PKS0858-004, this implies that the effective 1.4 GHz to 500 $\mu$m spectral index has to be steeper than 0.3, i.e., steeper than the effective indices of the two detected blazars (0.13 and 0.25), but still within the range observed for these sources.

Based on the best-fit model estimate of the synchrotron peak frequency, $\nu_p$, both [HB89] 0906+015 and PKS 0907+022 fall in the low frequency synchrotron peaked (LSP) category. This suggests that the H-ATLAS sample will set strong constraints on the abundance of blazars with low ($\lesssim 10^{13}$ Hz) values of $\nu_p$, which, in the blazar sequence scenario, are those with more powerful jets, more luminous accretion disks and higher black hole masses (Ghisellini et al. 2010a).

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